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Conditions for condensation and preservation of amorphous ice and crystallinity of astrophysical ices

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Abstract. Conditions for formation and preservation of amorphous ice formed through condensation of water vapor on a substrate is investigated both theoretically and experimentally. The kinetic consideration of deposition of vapor leads to the condition for formation of amorphous ice on a cold substrate: the flux of water vapor onto the substrate should be larger than a critical flux \( D_s/a^4 = F_c \), where \( D_s \) is the surface diffusion coefficient of water molecules on the substrate and \( a \) the lattice constant of the substrate. The validity of the derived condition has been confirmed by measuring the critical flux of water vapor onto a substrate of polycrystalline cubic ice as a function of temperature. From the measured \( F_c \), a surface diffusion coefficient of H2O molecules on polycrystalline cubic ice has been determined. With regard to the preservation condition, we derive a theoretical expression of the time scale \( t_c \) for crystallization of amorphous ice, which has the same form at low temperatures as the empirical formula used so far. A self-diffusion coefficient of amorphous ice is estimated from a comparison between the theoretical and empirical formulas of \( t_c \). The results are applied to condensation of ice in molecular clouds, circumstellar envelopes of late-type stars, and the primordial solar nebula. Discussion is given on the crystallinity of ices in these sites.

Key words: molecular processes – solar system: formation – stars: circumstellar matter – interstellar medium: clouds

1. Introduction

Ice is observed in various astrophysical sites (see Klinger 1983), and is a major component in the low-temperature environments. Understanding of ice condensation provides a basis for the study of the formation and evolution of interstellar and solar-system icy grains. So far, the equilibrium condensation theory (e.g., Lewis 1972; Yamamoto et al. 1983) has been widely applied to estimate condensation temperatures of ice in the cosmic environments because of its simplicity. On the other hand, it has been pointed out theoretically that kinetic effects in condensation plays an essential role in determining grain characteristics such as grain size, their number density, and condensation temperatures (Yamamoto & Hasegawa 1977; Draine & Salpeter 1977; Seki & Hasegawa 1983) as well as crystalline structure (Seki & Hasegawa 1981; Gail & Sedlmayr 1984). Recently it has been pointed out that revealing whether deposited ice is amorphous or crystalline provides a key clue to the origin and thermal evolution of comets (Yamamoto 1991).

While there are many polymorphs of H2O ice, phase relations at low temperatures and low pressures are especially relevant to discussing condensation of ice in space. The experimentally determined structure of ices formed by deposition of water vapor onto substrates at low temperatures is compiled by Hobbs (1974) as a function of temperature. According to his compilation, hexagonal ice (ice Ih) is formed as a stable phase above 170 K, metastable cubic ice (ice Ic) at temperatures between 170 and 130 K, and amorphous ice below 130 K. Most researchers presumed the transition temperature between ice Ic and amorphous ice to be 130 K (e.g., Seki & Hasegawa 1983). It should be pointed out, however, that the transition temperatures cannot be determined uniquely but have been shown experimentally to be highly dependent on the condensation rate (Olander & Rice 1972; Kouchi 1987) and pressure as well. On the other hand, conditions for formation of amorphous Ge studied experimentally (Sloope & Tiller 1962; Krikorian & Sneed 1966) are that the growth rate should be larger than a critical value and that the substrate temperature should be lower than a critical temperature. Although there are several qualitative analyses for Ge of the effect of the growth rate on epitaxy (Krikorian & Sneed 1966) and of crystal growth in amorphous phase (Narayan 1982), these analyses are not enough to be applied to formation of amorphous solids in general, especially to amorphous ice, because they studied the growth of Ge only on the single crystal

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substrate. It is therefore highly desirable to clarify the formation conditions of amorphous solids that can be applicable to a wide variety of formation environments including astrophysical ones.

In the present study, we shall investigate the formation and preservation conditions of amorphous solids with particular attention to amorphous ice. On the basis of kinetics of condensation and crystal growth, we shall derive criteria for formation of amorphous solids and its preservation in Sect. 2 together with results of the experiment of condensation of amorphous ice. Comparison between the theoretical and empirical formulas for crystallization time scale is made in Sect. 3. We shall give a physical basis of the empirical formula, and evaluate a self-diffusion coefficient of amorphous ice. Crystallinity of ices in astrophysical sites is discussed in Sect. 4.

2. Formation conditions of amorphous ice

For obtaining amorphous ice the following two conditions should be satisfied: (1) formation of amorphous ice at the time of condensation, and (2) its preservation against crystallization in the relevant environment. We shall discuss the two processes separately first and combine the conditions later.

2.1. The formation condition at the time of condensation

Condensation of ice usually occurs via vapor deposition on a cold substrate. Water vapor is usually deposited on a cooled substrate in a vacuum chamber in a laboratory (Hobbs 1974) and on the grain surfaces in cosmic conditions (Seki & Hasegawa 1983). In these cases the substrate is so cold that sublimation of adsorbed vapor molecules, which are called adatoms hereafter, is negligible. Thus we assume that the sublimation time scale \( \tau_s \) is much longer than the collision mean free time \( \tau_c \) of adatoms diffusing on a surface of the substrate (Irisawa et al. 1990):

\[
\tau_s \gg \tau_c. \tag{1}
\]

Amorphous ice forms if the adatoms are "frozen" in the vicinity of each of the adsorption points and cannot settle to the crystalline configuration. Namely, amorphous ice may be formed under the condition that the diffusion distance during the time of coverage of the surface by adatoms is smaller than the lattice constant \( a \) of crystalline ice. For the impinging flux of the water vapor \( F \), one has the coverage time \( t_{\text{cover}} \) as:

\[
t_{\text{cover}} = 1/Fa^2. \tag{2}
\]

The diffusion distance during \( t_{\text{cover}} \) is given by \( \sqrt{Ds t_{\text{cover}}} \), where \( D_s \) is the surface diffusion coefficient. Thus the formation condition of amorphous ice is expressed by

\[
D_s t_{\text{cover}} < a^2, \tag{3}
\]

which may be rewritten in terms of the flux \( F \) as

\[
F > D_s/a^4 \equiv F_c, \tag{4}
\]

which is equivalent to the condition obtained by Gail & Sedlmayr (1984).

To check the criterion (4), we made the following experiment. First a thin film of polycrystalline ice \( I_e \), 100 to 200 Å in thickness, was prepared at 140 K by deposition of water vapor onto a metal substrate in a vacuum chamber at a pressure of \( 5 \times 10^{-9} \) Torr by the method described previously (Kouchi 1990; Kouchi & Kuroda 1990; Kouchi et al. 1992). During an additional deposition of water vapor at a constant flux \( F \), the substrate was cooled slowly at the rate of 0.1–1 K/min, and structural changes of the ice surface were observed in situ by the reflection electron diffraction method using 20 keV electrons. The deposition rates were measured from separate calibration experiments by using the laser interferometry method. For varying fluxes a transition from crystalline to amorphous ice was observed to occur at varying temperatures. From these data we obtained the dependence of the critical flux \( F_c \) on the temperature \( T \).

Figure 1 shows an example of reflection electron diffraction patterns of the surface of the ice film. In the high temperature region, the deposited ice was polycrystalline ice \( I_e \) (Fig. 1a). With decreasing substrate temperature amorphous ice began to condense at a critical temperature \( T \) (Fig. 1b). Further decrease of the substrate temperature leads to continuation of condensation of amorphous ice. These observations clearly indicate the presence of a critical flux from condensation of ice \( I_e \) to amorphous ice for a given temperature \( T \) as was predicted.

Figure 2 shows the critical flux \( F_c \) experimentally determined as a function of \( 1/T \), where the straight line shows a least-squares fit. The linear relation between \( \ln F_c \) and \( 1/T \) is consistent with the theoretical prediction that \( F_c \) is proportional to the surface diffusion coefficient \( D_s \) [Eq. (4)], which is expressed by

\[
D_s = D_{s0} \exp (-E_s/kT). \tag{5}
\]

From the measured \( F_c(T) \) shown in Fig. 2, we may determine, with the use of Eq. (4), \( D_{s0} \) and \( E_s \), and thus the surface diffusion coefficient \( D_s \) of \( H_2O \) molecules on polycrystalline ice \( I_e \). The result is \( E_s/k = 4590 \) K and \( D_{s0} = 1.74 \times 10^5 \) cm²/s for \( a = 4.5 \) Å. This surface diffusion coefficient is regarded as an upper limit, since the criterion (4) is a sufficient condition for forming amorphous ice. As will be seen in Fig. 4, \( D_s \) on polycrystalline ice \( I_e \) is much smaller than that on the (0001) face of ice \( I_b \) single crystal. This indicates that the surface of polycrystalline ice \( I_e \) is very rough on a molecular scale, which is also supported by the fact that the reflection electron diffraction pattern (Fig. 1a) is the same as the transmission one.

2.2. Formation of amorphous ice by slow vapor deposition in the laboratory

The formation condition (4) for amorphous ice that the deposition rate be larger than a critical one \( F_c \) appears at first to be completely opposite to the condition stated in Olander and Rice (1972) and confirmed by experiments performed to analyse the structure of amorphous ice (Narten et al. 1976) as well as by experiment designed to study absorption spectra (Breuckers 1991). In those experiments, increasing rate of deposition
Fig. 1 a and b. Reflection electron diffraction patterns of a ice Ic deposited at 100 K and b amorphous ice deposited at 99.3 K on the surface of polycrystalline ice Ic. The deposition rate is $2 \times 10^{-8}$ cm s$^{-1}$.

was accompanied by less and less amorphicity and ultimately to crystallinity (Kouchi et al. 1992a,b). The explanation for this apparent discrepancy with condition (4) is simply described.

The water molecule deposition is always accompanied by the release of the latent heat of fusion. With relatively thick samples the thermal conductivity of the initially amorphous ice is so small that localized heating cannot be transported away rapidly enough to prevent diffusion accompanying higher deposition rates and conversely. In case the sample is unshielded from the ambient room-temperature radiation field there is even less possibility for removing the heat from the local sites. As already stated the essential of condition (4) is that H$_2$O molecules may diffuse so that if local heating with increased flux enhances the diffusion this leads to crystallinity. A schematic of the variability of the systems leading first to decreasing and then increasing crystallinity with the flux is shown in Fig. 3 with the different regimes clearly separated. The experiments described in Sect. 2.1 lead to the condition for amorphicity with increasing flux and follow the left side of the curve in Fig. 3. The experiments described in Kouchi et al. (1992a, b) follow the right portion of the curve in Fig. 3.

2.3. Condition for preservation of amorphous ice

We shall derive the time scale for crystallization of amorphous ice on the basis of the theories of nucleation and crystalline growth. If the time scale of crystallization is longer than the relevant time scale of the system concerned, amorphous ice will be preserved in practice.

Crystalline ice is formed through nucleation of crystalline clusters in amorphous phase and subsequent growth of the clusters. The volume fraction $\theta$ of crystalline ice at a time $t$ depends on the nucleation rate $J$ and the growth rate $w$, both of which are functions of the temperature. For constant $J$ and $w$, the crystalline fraction is given from the kinetic theory of crystallization formulated by Kolmogorov (see Hobbs 1974) by

$$\theta(t) = 1 - \exp(-\alpha J w^3 t^4),$$

(5)

where $\alpha$ is a geometrical factor depending on the morphology of the crystal growth, and is given by $\alpha = 2$ for cubic growth and $\alpha = \pi/3$ for spherical growth. We define the crystallization time scale $t_c$ by

$$t_c = (\alpha J w^3)^{-1/4}.$$  

(6)

To compare $t_c$ with the experimental results, we shall rewrite Eq. (6) in what follows.

The change in the Gibbs free energy $\Delta G$ for formation of a crystalline cluster of radius $r$ in an amorphous phase is expressed by

$$\Delta G = -\frac{4\pi r^3}{3\Omega} \Delta G_v + \frac{4\pi r^2}{\Omega^{2/3}} \sigma,$$

(7)

where $\Omega$ is volume of a water molecule, $\sigma = \gamma \Omega^{2/3}$ with $\gamma$ being the interfacial tension, and $\Delta G_v$ is the free energy difference.
from $\frac{d\Delta G}{dr} = 0$, and the free energy difference for $r = r^*$ is given by

$$\Delta G^* = \frac{16\pi\sigma^3}{3\Delta G_v^2}.$$  \hfill (10)

The nucleation rate $J$ of crystalline nuclei in amorphous phase at temperature $T$ may be written (Seki & Hasegawa 1981) as

$$J = Z \frac{4\pi r^{*2}}{\Omega^{2/3}} \frac{D}{\Omega^{2/3}} \exp \left( -\frac{\Delta G^*}{kT} \right),$$  \hfill (11)

where $Z = \Omega^{2/3}(\sigma/kT)^{1/2}/2r^{*2}$ is the Zeldovich factor, $D$ is the self-diffusion coefficient, and $k$ is the Boltzmann constant. With the use of Eqs. (8) and (10), $J$ is expressed by

$$J = \frac{2\pi}{\Omega} \left( \frac{\sigma}{kT} \right)^{1/2} \frac{D_0}{\Omega^{2/3}} \exp \left[ -\frac{1}{kT} \left( \frac{E_a + 16\pi\sigma^3}{3L^2} \right) \times \left( \frac{T_m}{T_m - T} \right)^2 \right],$$  \hfill (12)

where we put

$$D = D_0 \exp \left( -\frac{E_a}{kT} \right)$$  \hfill (13)

with $E_a$ being the activation energy of self-diffusion.

The growth rate of crystalline clusters is given (Seki & Hasegawa 1981) by

$$w = \frac{D}{\Omega^{1/3}} \left[ 1 - \exp \left( -\frac{\Delta G_v}{kT} \right) \right],$$  \hfill (14)

where the second term in the bracket indicates the effect of remelting of the crystalline surface, and may be ignored at the low temperatures of concern here. Then $w$ is approximated by

$$w = \frac{D_0}{\Omega^{1/3}} \exp \left( -\frac{E_a}{kT} \right).$$  \hfill (15)

Substituting Eqs. (12) and (15) into Eq. (6), we obtain the crystallization time scale as a function of $T$:

$$t_c = \left( \frac{1}{2\pi\alpha} \right)^{1/4} \left( \frac{kT}{\sigma} \right)^{1/8} \frac{\Omega^{2/3}}{D_0} \exp \left[ -\frac{1}{kT} \left( \frac{E_a + 4\pi\sigma^3}{3L} \right) \times \left( \frac{T_m}{T_m - T} \right)^2 \right].$$  \hfill (16)

The condition for preservation of amorphous ice is that $t_c$ is longer than the relevant time scale of the system. In other words, given the time scale $t_c$, a typical temperature $T$ of the relevant system should be lower than the temperature $T_c(t_c)$:

$$T < T_c(t_c),$$  \hfill (17)

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The importance of knowing the self-diffusion coefficient of amorphous ice has been invoked for understanding a sublimation process of CO from amorphous H$_2$O-CO ice (Sandford & Allamandola 1988) and for the theoretical analysis of radiation-induced amorphization of ice crystals (Kouchi & Kuroda 1991). There has, however, been no direct measurement or estimation. We shall show that the crystallization time scale derived in Sect. 2.2 provides a method to obtain the self-diffusion coefficient $D$, and estimate $D$ of amorphous ice.

Experimentally the crystallization time scale of amorphous ice has been measured as a function of the temperature ($115 \lesssim T \lesssim 145$ K) by X-ray diffraction (Dowell & Rinfret 1960), electron diffraction (Dubochet et al. 1982), and infrared spectroscopy (Hardin & Harvey 1973; Schmitt et al. 1989). The time scale of crystallization is expressed empirically by

$$t_c = A \exp(E/kT),$$

where $A$ and $E$ are constants determined from the experiments. Here we note that $t_c$ given by Eq. (16) reduces to the same
form as the experimental formula of Eq. (19) when \( T \ll T_m \). Comparison of Eq. (16) for \( T \ll T_m \) with Eq. (19) yields

\[
E = E_a + \frac{4 \pi \sigma^2}{3L^2}
\]

\[
A = \left( \frac{1}{2\pi \alpha} \right)^{1/4} \left( \frac{kT_{\text{exp}}}{\sigma} \right)^{1/8} \frac{\Omega^{2/3}}{D_0},
\]

where \( T_{\text{exp}} \) is a typical temperature in the measurements. Uncertainty in the choice of \( T_{\text{exp}} \) does not much affect the value of \( A \) because of the weak dependence of \( A \) on \( T_{\text{exp}} \).

The relations (20) and (21) imply that the self-diffusion coefficient \( D \) expressed by Eq. (13) may be determined from measured \( A \) and \( E \). For numerical estimates we take \( A = 9.54 \times 10^{-14} \) and \( Ea/k = 5370 \) K (Schmitt et al. 1989). Other quantities that we take are \( \Omega = 3.25 \times 10^{-23} \) cm\(^3\) (Yokoyama & Kuroda 1990), \( L = 2.6 \times 10^{-14} \) erg/molecule (Ghormley 1968), and \( \sigma = 1.0 \times 10^{-14} \) erg (Fletcher 1970). Then we have \( E_a/k = 5370 \) K and \( D_0 = 6.1 \times 10^{-3} \) cm\(^2\) s\(^{-1}\) for the transition to cubic ice (\( \alpha = 2 \)). It is interesting to note that the \( E_a \)-value obtained agrees with that expected from the van Liiempt relation \( E_a/k = 17T_m \) (e.g. Poirier 1991) for \( T_m = 273 \) K within a factor of 1.15.

Figure 5 compares the self-diffusion coefficient \( D \) of H\(_2\)O molecules in amorphous ice obtained here with those in ice I\(_h\) (Goto et al. 1986) and in liquid water (Wang et al. 1953) (the lines a, b, c). It is clearly seen that \( D \) of amorphous ice is larger than that of ice I\(_h\) by one to two orders of magnitude at the temperatures where amorphous ice is preserved. This suggests that, compared with ice I\(_h\), amorphous ice has porous structure in a molecular cluster scale, which is also suggested by its very low thermal conductivity (Kouchi et al. 1992) and large surface area (Schmitt et al. 1987; Mayer & Pletzer 1986; Zhang & Buch 1990; Bar-Nun et al. 1985). However the porosity appears to be limited to about \( P = 0.1 \) as derived from the infrared absorption strength of H\(_2\)O (Breukers 1991). It should be pointed out that the structure of amorphous ice not only depends on the deposition rate and deposition temperature (Olander & Rice 1972; Kouchi 1987; Zhang & Buch 1990; Narten et al. 1976) but also changes with increasing temperature. Schmitt et al. (1989) measured the crystallization time scale \( t_c \) of amorphous ices at temperatures between 125 and 140 K, which were deposited at 10 K. With increasing temperature, the structure of amorphous ice deposited at 10 K changes to another structure at around 90 and 135 K (Kouchi 1990; Laufer et al. 1987). Therefore, the self-diffusion coefficient \( D \) estimated in the present study is regarded as that of the high temperature form of amorphous ice.

Figure 5 also shows that the surface diffusion coefficients \( D_s \) of H\(_2\)O molecules on the surfaces of (0001) of ice I\(_h\) (d) and of polycrystalline ice I\(_c\) (e). One sees that \( D_s \) on polycrystalline ice is much larger than \( D_s \) on amorphous ice. Since the surface of amorphous ices is rougher than that of polycrystalline ice microscopically (Schmitt et al. 1987; Mayer & Pletzer 1986; Zhang & Buch 1990; Bar-Nun et al. 1985; Laufer et al. 1987), it is expected that the surface diffusion coefficient \( D_s \) on amorphous ice should be between \( D_s \) on polycrystalline ice I\(_c\) and \( D \) in amorphous ice.

4. Crystallinity of astrophysical ices

Crystallinity of astrophysical ice provides a clue to investigating the evolution of grains in interstellar space and the origin and evolution of the solar-system icy bodies such as cometary nuclei. In this section we briefly discuss crystallinity of astrophysical ices on the basis of the criteria described in Sect. 2.

4.1. Circumstellar envelopes

Crystalline H\(_2\)O ice is identified from the distinct emission features at 44 and 60 \( \mu \)m in the circumstellar envelope of the evolved bipolar nebula IRAS 09371+1212 (Omont et al. 1990). The observation suggests the possibility of condensation of crystalline ice in circumstellar envelopes of evolved stars, although the evolutionary stage of the star is still uncertain (Robinson et al. 1992).

In circumstellar envelopes of evolved stars, water ice condenses to form icy mantles on the surfaces of refractory grains already condensed as the gas cools down with expansion. For an estimate of the impinging flux of H\(_2\)O molecules, we assume a simple model that the circumstellar envelope is spherically symmetric with constant outflow velocity \( v \) and that the temperature of the gas at the distance \( r \) from the central star is expressed by

\[ T(r) = T_{\text{phot}}(r/r^{\alpha})^{\beta}, \]

where the asterisk represents the values at the photosphere of the central star and \( \beta \) is a parameter depend-
ing on the cooling rate of the gas. Then the mass density \( \rho \) at distance \( r \) from the central star with mass loss rate \( \dot{M} \) is given by

\[
\rho = \frac{\dot{M}}{4\pi r^2 v} = \left( \frac{\dot{M}}{L} \right) \frac{\sigma_T T^4}{v} \left( \frac{T}{T_*} \right)^{2/\beta},
\]

where \( T_* \) is the effective temperature of the central star, \( L = 4\pi r^2 \sigma_T T^4 \) its luminosity, and \( \sigma_T \) the Stefan-Boltzmann constant. In an envelope of an oxygen-rich star all carbon atoms are locked into CO molecules and the remaining O atoms form H\(_2\)O molecules (Jura \& Morris 1985 and references therein). The concentration of H\(_2\)O molecules \( n_{\text{H}_2\text{O}} \) is expressed by \( n_{\text{H}_2\text{O}} = f_{\text{H}_2\text{O}} \rho / \mu m_H \), where \( \mu \) is the mean molecular weight of the gas, \( f_{\text{H}_2\text{O}} \) the number fraction of H\(_2\)O molecules, and \( m_H \) mass of an hydrogen atom. The impinging flux of H\(_2\)O molecules per unit surface area of the grains is given as a function of the gas temperature \( T \) by

\[
F_{\text{H}_2\text{O}} = \frac{f_{\text{H}_2\text{O}} \rho}{\mu m_H} \left( \frac{kT}{2\pi \mu m_H} \right)^{1/2}
= 2.4 \times 10^4 \left( \frac{2}{\mu} \right) \left( \frac{\dot{M}}{10^{-5} M_\odot \text{yr}^{-1}} \right) \left( \frac{10^4 L_\odot}{L} \right) \left( \frac{f_{\text{H}_2\text{O}}}{10^{-4}} \right)
\times \left( \frac{10 \text{ km s}^{-1}}{v} \right) \left( \frac{T}{3000 \text{ K}} \right)^{4-2/\beta}
\times \left( \frac{T}{100 \text{ K}} \right)^{2/\beta+1/2} \text{ cm}^{-2}\text{s}^{-1},
\]

where \( \mu_{\text{H}_2\text{O}} = 18 \) is the molecular weight of an H\(_2\)O molecule. We put \( \beta = 0.5 \) in Eq. (23) for the numerical estimate. Calculations show that condensation of H\(_2\)O ice on silicate grains occurs at temperatures around 100 K for the parameters \( L \) ranging from \( 5 \times 10^3 \) to \( 10^4 L_\odot \) and \( M \) ranging from \( 10^{-6} \) to \( 10^{-4} M_\odot \) yr\(^{-1}\). The difference between the equilibrium temperature and the condensation temperature is at most 30 K in the range of the calculations.

Thus the impinging flux of H\(_2\)O molecules on substrate grains in circumstellar envelopes is \( F \sim 10^8 \text{ cm}^{-2}\text{s}^{-1} \) at \( T = 100 \text{ K} \), and is much lower than the critical flux for formation of amorphous ice \( F_c = 10^{22} \text{ cm}^{-2}\text{s}^{-1} \) as shown in Fig. 6. In consequence icy mantles formed in circumstellar envelopes of evolved stars is crystalline. Here we should keep in mind that the optical thickness of the circumstellar envelopes must be larger than a certain value in order for crystalline ice to be formed and preserved; otherwise interstellar UV photons penetrating into the envelope could dissociate H\(_2\)O molecules (Jura \& Morris 1985), and amorphize the crystalline H\(_2\)O ice already formed (Kouchi \& Kuroda 1990). In any case the kinetic consideration for the formation of amorphous ice leads to the conclusion that H\(_2\)O ice should be crystalline at the time of condensation in circumstellar envelopes of evolved stars. Note that our discussion is different from the previous discussions by Omont et al. (1990) and Robinson et al. (1992): They consider that crystalline ice is formed by crystallization of the amorphous ice mantle after the condensation, whereas our conclusion is that crystalline ice is formed at the condensation because of \( F < F_c \). Our conclusion is consistent with the experimental results obtained by studying the 45 \( \mu \text{m} \) absorption feature of ice and using it in the modelling of silicate-core ice mantle grains in the envelope of IRAS 09371+22 (Breuker 1991).

### 4.2. Molecular clouds

Water ice observed in molecular clouds is in an amorphous phase, which is concluded from the broad absorption feature around 3.1 \( \mu \text{m} \) in comparison with the spectral feature of crystalline H\(_2\)O ice (e.g. van de Bult et al. 1985; Hagen et al. 1981; Whittet 1992).

In molecular clouds the temperature of grains is so low (typically \( \sim 10 \text{ K} \)) that most molecules as well as atoms which strike their surface will stick. The formation of H\(_2\)O molecules as a mantle on the grain occurs primarily as a result of sequential surface reactions of O + H \( \rightarrow \) OH and OH + H \( \rightarrow \) H\(_2\)O. The rate at which this occurs can be described as a flux of H\(_2\)O molecules derived from sticking of O atoms:

\[
F_{\text{H}_2\text{O}} = n_{\text{H}_2} f_0 \left( \frac{kT}{2\pi \mu m_H} \right)^{1/2}
= 2.7 \times 10^3 \left( \frac{n_{\text{H}_2}}{10^4 \text{ cm}^{-3}} \right) \left( \frac{f_0}{10^{-4}} \right)
\times \left( \frac{T}{10 \text{ K}} \right)^{1/2} \text{ cm}^{-2}\text{s}^{-1}.
\]
However, at each stage of surface reactions an energy release on the order of an eV occurs which locally heats the surface. This will lead to a possible displacement of the H2O molecule formation point and consequently so does the surface diffusion formed. An upper limit of the displacement is estimated in the Appendix, in which the surface diffusion distance induced by heating due to molecule formation is calculated. The result on the order of an eV occurs which locally heats the surface.}

4.3. The primordial solar nebula

To discuss crystallinity of ice in the primordial solar nebula, one has to know the temperature history during formation of the solar nebula. But the details are still uncertain. In view of the present theory of formation of the primordial solar nebula (e.g. Hayashi et al. 1985), we simplify the temperature history to the two stages as schematically shown in Fig. 7: (1) Heating followed by cooling at the stage of the solar nebula formation. The heating is caused by shock waves induced by infalling of the gas in the parent molecular cloud onto the nebular disk surfaces. The cooling is due to thermal emission from the nebular surfaces. (2) A steady stage, which is reached after the cooling of the solar nebula. We shall estimate (a) the maximum temperature for preservation of amorphous ice in the parent molecular cloud at stage (1), (b) the region in the solar nebula where amorphous ice is preserved, and (c) crystallinity of the ice recondensed in the cooling of the solar nebula.

First we estimate (a) the temperature for preservation of amorphous ice in the parent molecular cloud. The ice in the parent molecular cloud could crystallize by heating at stage (1). The crystalline fraction \( \theta \) given by Eq. (5) is that for a constant temperature. When the temperature varies with time, the change in the nucleation rate \( J \) and the growth rate \( w \) should be taken into account, which leads to a rather complicated formula. Instead we adopt a phenomenological approach and write down the equation describing the rate of the change in the crystallization degree (Haruyama et al. 1993) as

\[
\frac{\partial \theta}{\partial t} = \frac{1 - \theta}{t_c} e^{-E/kT} - \frac{(1 - \theta)e^{-E/kT}}{A}, \tag{25}
\]

where \( t_c \) is the crystallization time scale given by Eq. (19). Integration of Eq. (25) leads to

\[
1 - \theta = \exp \left( -\frac{1}{A} \int_0^t e^{-E/kT} dt \right), \tag{26}
\]

where \( T \) is the maximum temperature reached is higher than \( T_{\text{max}} \) determined by

\[
1 - \theta = \exp \left[ -\frac{2\pi k}{E(-\partial^2 T/\partial t^2)_{\text{max}}} \right] \tag{27}
\]

Substantial crystallization occurs when the maximum temperature reached is higher than \( T_{\text{max}} \) determined by

\[
1 - \theta = \exp \left[ -\frac{2\pi k}{E(-\partial^2 T/\partial t^2)_{\text{max}}} \right] \tag{28}
\]

Since the crystallization rate \( 1/t_c = A^{-1} e^{-E/kT} \) is a rapidly increasing function of \( T \), the crystallization proceeds most efficiently near the maximum temperature \( T_{\text{max}} \) in the course of heating of the solar nebula. In other words the integrand has a sharp peak at \( t = t_{\text{max}} \), the time of the maximum temperature \( T_{\text{max}} \). The integration may thus be carried out with the use of the conventional saddle-point method, and one obtains

\[
1 - \theta = \exp \left[ -\frac{2\pi k}{E(-\partial^2 T/\partial t^2)_{\text{max}}} \right] \tag{27}
\]

We evaluate \( (-\partial^2 T/\partial t^2)_{\text{max}} \) in Eq. (28) roughly to be \( \Delta T/t_{\text{heat}}^2 \), where \( \Delta T \approx 100 \text{ K} \) is the temperature elevation and \( t_{\text{heat}} \) is the heating time scale. According to a model of the solar nebula formation (e.g. Hayashi et al. 1985), the parent molecular cloud collapses isothermally up to the gas density of \( \sim 10^{-13} \text{ g cm}^{-3} \). Afterward the temperature of the gas increases with increasing density of the cloud. A shortest estimate of \( t_{\text{heat}} \) may be taken as the free-fall time from the gas density \( \sim 10^{-13} \text{ g cm}^{-3} \), which gives \( t_{\text{heat}} \sim 10^{10} \text{ s} \). An upper limit of the time scale, on the other hand, may be taken as the lifetime of the molecular cloud, i.e. \( t_{\text{heat}} \sim 10^{15} \text{ s} \). Thus \( (-\partial^2 T/\partial t^2)_{\text{max}} \) would be in the interval of \( 10^{-8} \text{ to } 10^{-13} \text{ K s}^{-2} \).

The temperature \( T_{\text{max}} \) is calculated for a wide range of \( (-\partial^2 T/\partial t^2)_{\text{max}} \) of \( 10^{-8} \text{ to } 10^{-13} \text{ K s}^{-2} \). It is found that \( T_{\text{max}} \) is insensitive to \( (-\partial^2 T/\partial t^2)_{\text{max}} \) and is 110 to 120 K. Note that these temperatures are rather close to the sublimation temperature of ice, implying that amorphous ice in the molecular cloud is preserved so long as it survived without sublimation.

The amorphous ice which survived during the formation of the solar nebula enters the stage of the steady-state solar nebula (see Fig. 7), in which crystallization proceeds at a rate depending
on the temperature where the ice is placed. We discuss (b) the region in the solar nebula where amorphous ice is preserved during the lifetime of the solar nebula. To get a quantitative picture, one requires the solar nebula model which gives the temperature distribution. We adopt the steady nebula model by Hayashi (1981), in which the temperature at heliocentric distance \( r \) is given by

\[
T(r) = 280 (\text{AU} / r)^{1/2} \text{K}.
\]

The crystallization time scale calculated from Eq. (19) together with Eq. (29) is shown as a function of the heliocentric distance \( r \) in Fig. 8. It is seen that amorphous ice is preserved in the region where \( r > 12 \text{ AU} \); i.e., outside the Saturnian region, during the nebular lifetime of \( \sim 10^8 \text{ yr} \). This implies that icy grains which formed the Uranian and Neptunian satellites and comets were initially amorphous, if they were formed from the icy grains preserved from the molecular cloud stage.

Finally we discuss (c) the crystallinity of ice recondensed in the solar nebula. The present theories of the solar nebula formation have not clarified which region in the solar nebula was heated up to the temperature above the sublimation temperature of ice. One may expect that the temperature in the inner region of the nebula rose enough at stage (1) that ice in the parent cloud sublimed and did not recondense even at stage (2). The region where ice could exist is limited to the region beyond the asteroidal belt (Hayashi 1981). There is an alternative possibility that ice in the parent molecular cloud first sublimed during the formation of the solar nebula and subsequently recondensed when the solar nebula cooled down (see Fig. 7), in contrast with the case of the preservation of ice in the parent cloud discussed above. The ice condenses on the surface of the pre-existing refractory grains as in the case of circumstellar envelopes. The impinging flux of \( \text{H}_2\text{O} \) molecules is given by

\[
F_{\text{H}_2\text{O}} = n_{\text{H}_2}(r) f_{\text{H}_2\text{O}} \left( \frac{kT}{2\pi m_{\text{H}_2\text{O}} n_{\text{H}_2}} \right)^{1/2}
= 6.4 \times 10^{14} \left( \frac{r}{\text{AU}} \right)^{-3} \left( \frac{f_{\text{H}_2\text{O}}}{10^{-4}} \right) \text{cm}^{-2} \text{s}^{-1},
\]

where the gas density \( n_{\text{H}_2}(r) \) as a function of the heliocentric distance \( r \) is taken from the Hayashi model (1981), which gives \( n_{\text{H}_2}(r) = 4.4 \times 10^{14} (r/\text{AU})^{-1.1/4} \text{cm}^{-3} \) at the midplane of the solar nebula. This flux is compared with the critical flux \( F_c \) given by the criterion (4). The result is shown in Fig. 8, which indicates that crystalline ice forms in the whole region of the solar nebula in contrast with the case of the preservation of amorphous ice in the parent molecular cloud.

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Appendix: upper limit of the diffusion distance of an \( \text{H}_2\text{O} \) molecule formed on the grain surface

Consider an \( \text{H}_2\text{O} \) molecule formed at \( t = 0, r = 0 \) on the surface, where energy \( E \) is released instantaneously. We assume that the energy \( E \) is transferred to the grain by conduction, and approximate the grain surface by a plane. The temperature at distance \( r \) from the formation point at time \( t \) is then given by a solution of the equation of heat conduction as

\[
T(r, t) = \frac{E}{4\rho c_p (\pi \chi t)^{3/2}} \exp \left( -\frac{r^2}{4\chi t} \right),
\]

where \( \rho \) is the density of the grain material, \( c_p \) is the specific heat per unit mass, and \( \chi \) is the thermal diffusivity, and \( \rho c_p \) and \( \chi \) are assumed to be constant. From Eq. (A1) one sees that at a given distance \( r \) the temperature reaches a maximum of

\[
T_{\text{max}} = \left( \frac{6}{\pi e} \right)^{3/2} \frac{E}{4\rho c_p r^3}
\]

at a time \( t_{\text{max}} = r^2/6\chi \). Furthermore it should be kept in mind that the temperature \( T(r_1, t) \) is always higher than the temperature \( T(r_2, t) \) for \( r_1 < r_2 \).

For accurate evaluation of the surface diffusion distance of a molecule formed at \( r = 0 \) with the energy release \( E \), one has to solve the diffusion equation where the diffusion coefficient,
which is a function of temperature given by Eq. (A1), varies with time \( t \) and distance \( r \). Here we estimate an upper limit of the diffusion distance in a simple manner, which is sufficient for the present purpose.

First we investigate the surface diffusion distance of an \( \text{H}_2\text{O} \) molecule initially placed at distance \( r \). An upper limit of the diffusion distance \( d_{\text{max}} \) may be written as

\[
d_{\text{max}}^2(r) = \int_0^\infty D_e dt = a_2^2 \nu \int_0^\infty e^{-E_a/kT(r,t)} dt. \tag{A3}
\]

The actual diffusion distance of a molecule placed at \( r \) at \( t = 0 \) is smaller than \( d_{\text{max}}(r) \) given by Eq. (A3), since the temperature becomes lower as the molecule diffuses to a point at \( r > r_\text{m} \) and thus the diffusion coefficient decreases compared with that at the initial point \( r \). Since the temperature \( T(r,t) \) given by Eq. (A1) has a sharp maximum \( T_{\text{max}} \) at time \( t_{\text{max}} \), the integration may be carried out by the saddle point method. The result is given by

\[
\frac{d_{\text{max}}^2}{a_2^2} = \frac{2\nu}{\chi} \left( \frac{\pi}{27K} \right)^{1/2} x^{1/2} e^{-Kx^2} \text{ with } x \equiv r/a \quad \text{and} \quad (A4)
\]

\[
K \equiv \frac{kT_{\text{max}}(r = a) = 4 \left( \frac{\pi \nu}{6} \right)^{1/2} \rho c_p E_a^3 k E}{k} = 3.87 \left( \frac{\rho}{g \text{ cm}^{-3}} \right) \left( \frac{c_p}{10^7 \text{ erg g}^{-1} \text{K}^{-1}} \right) \left( \frac{E_a}{10^3 \text{ K}} \right) \times \left( \frac{a}{4.5 \text{ Å}} \right)^3 \left( \frac{eV}{E} \right). \tag{A5}
\]

For representative values of \( K = 3.87, a = 4.5 \text{ Å}, \) and \( \nu = 10^{13} \text{ s}^{-1} \), and a lower limit of \( \chi = 8.9 \times 10^{-8} \text{ cm}^2 \text{ s}^{-1} \), which is the thermal diffusivity of amorphous ice (Kouchi et al. 1992; see also Haruyama et al. 1993), we have \( d_{\text{max}}^2/a_2^2 \approx 2 \times 10^{-9} \ll 1 \) for \( x = 2 \) and much smaller values for \( x > 2 \), indicating that molecules initially placed at \( x \geq 2 \) remain there even after the heat of formation has passed through. We denote by \( r_\text{m} \) the minimum distance where \( d_{\text{max}}^2/a_2^2 \) becomes less than unity. For the above estimate \( r_\text{m} = 2a \).

Next let us imagine the situation that a molecule formed at \( r = 0 \) and \( t = 0 \) diffuses to the distance \( r = r_\text{m} \) at a time \( t_\text{m} \). The diffusion distance of the molecule after the time \( t_\text{m} \) is less than \( a \) by the definition of \( r_\text{m} \). Thus a molecule formed at \( r = a \) cannot diffuse to the distance larger than \( r_\text{m} \). The \( r_\text{m} \) is less than 2\( a \) for the parameter values adopted above, and even if we take \( K \) as small as \( K = 1 \), \( r_\text{m} = 3a \). In conclusion the molecules formed are confined in the immediate vicinity of the formation place in general.

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